

ALMA Lensing Cluster Survey (ALCS): Overview and synergies with JWST

Kotaro Kohno^{*(1)},

(1) Institute of Astronomy, Graduate School of Science, The University of Tokyo, Tokyo, Japan; e-mail: kkohno@ioa.s.u-tokyo.ac.jp

The ALMA lensing cluster survey (ALCS) is a 96-hr large program to search for intrinsically faint continuum sources and line emitters with the assistance of gravitational lensing. All 33 cluster fields were selected from *HST/Spitzer* treasury programs including CLASH, Hubble Frontier Fields, and RELICS. The total sky area surveyed reaches $\sim 133 \text{ arcmin}^2$ down to a depth of $\sim 60 \mu\text{Jy beam}^{-1}$ (1σ) at 1.2 mm, yielding 141 secure blind detections of continuum sources and additional 39 sources aided by priors. Fig.1 shows rich ALMA 1.2 mm sources including known multiply imaged dusty galaxy at $z = 3.631$ and newly uncovered two sets of triple-image *H*-dropout sources [1, 2]. We find an association of an *H*-dropout ALMA source with a MUSE-selected galaxy group at $z = 4.32$ behind “El Gordo” [3]. Multi-wavelength mosaics across all 33 ALCS fields and photometric catalogs have been created [4], which have been exploited to conduct stacking analysis [5]. *Herschel* far-infrared and *Chandra* X-ray properties of ALCS sources have been investigated [1, 6]. An exceptionally magnified ($\mu \sim 20\text{--}160$), intrinsically faint (sub- L^* -class) low-mass ($M_* \sim 10^9 M_\odot$), bright [CII] line emitting star-forming galaxy at $z = 6.07$ [7, 8] is scheduled to be observed by *JWST/NIRCam*, *NIRSpec*, *VLT/MUSE*, and so on. *JWST* observations of lensing clusters including ALCS fields are now rapidly increasing, allowing us to characterize ALCS 1.2 mm continuum sources in e.g., Abell 2744 [9].

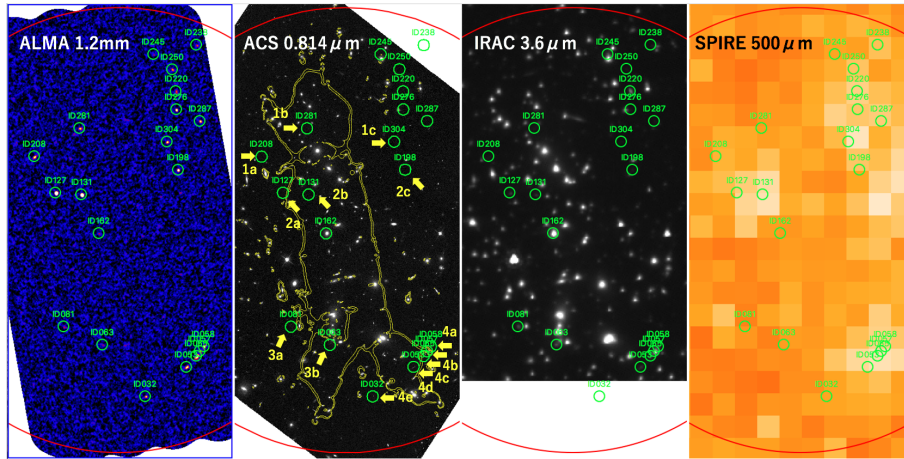


Figure 1. ALMA 1.2 mm continuum image of RXCJ0032.1+1808, along with the *HST/ACS*, *Spitzer/IRAC*, and *Herschel/SPIRE* images. Critical curves at $z = 4$ are shown in the *HST/ACS* panel, exhibiting the richness of multiple images of dusty star-forming galaxies, which are mostly invisible even in the *HST* near-infrared bands.

References

- [1] Sun, F.; Egami, E.; Fujimoto, S.; et al. 2022, *ApJ*, 932, 77
- [2] Tsujita, A., Kohno, K., Hatsukade, B., et al. URG1-GA 2023
- [3] Caputi, K. I.; Caminha, G. B.; Fujimoto, S. et al. 2021, *ApJ*, 908, 146
- [4] Kokorev, V.; Brammer, G.; Fujimoto, S. et al. 2022, *ApJS*, 263, 38
- [5] Jolly, J.-B.; Knudsen, K.; Laporte, N. et al. 2021, *A&A*, 652, A128
- [6] Uematsu, R.; Ueda, Y.; Kohno, K. et al. *ApJ*, in press (arXiv:2301.09275)
- [7] Laporte, N.; Zitrin, A.; Ellis, R. S. et al. 2021, *MNRAS*, 505, 4838
- [8] Fujimoto, S.; Oguri, M.; Brammer, G. et al. 2021, *ApJ*, 911, 99
- [9] Kokorev, V.; Jin, S.; Magdis, G. E. et al. submitted to *ApJL* (arXiv:2301.04158)